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Solar Neutrinos

M. Misiaszek

Jagiellonian University, Cracow

Symposium SFOF PTF "Standard Model and Beyond"
21-23.10.2022, Katowice

Talk overview

1. **Introduction and motivation for solar neutrino program**

- Direct probe of nuclear fusion
- Standard Solar Models: metallicity
- Neutrino oscillation parameters: solar sector (θ_{12}, Δm^2_{12})
- Survival probability P_{ee} as f(E_ν): matter effects, testing LMA-MSW prediction and its upturn

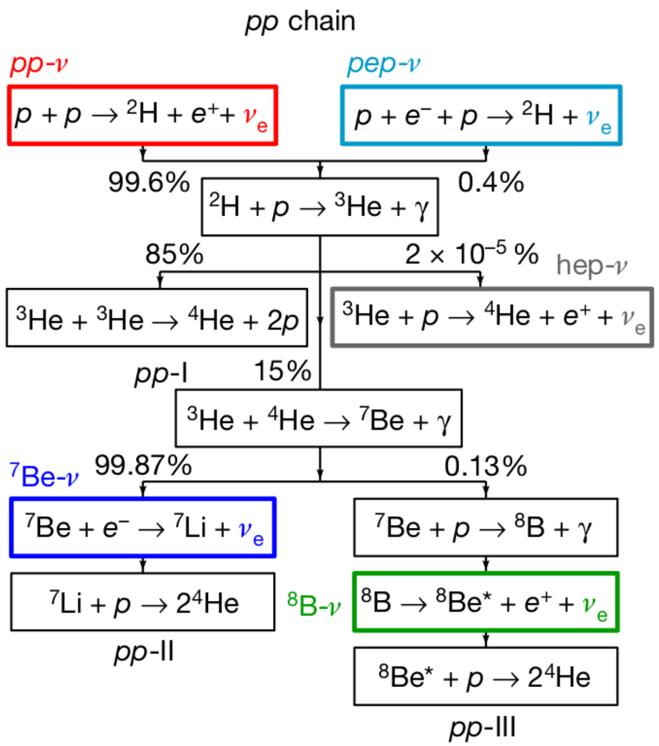
2. **Borexino**

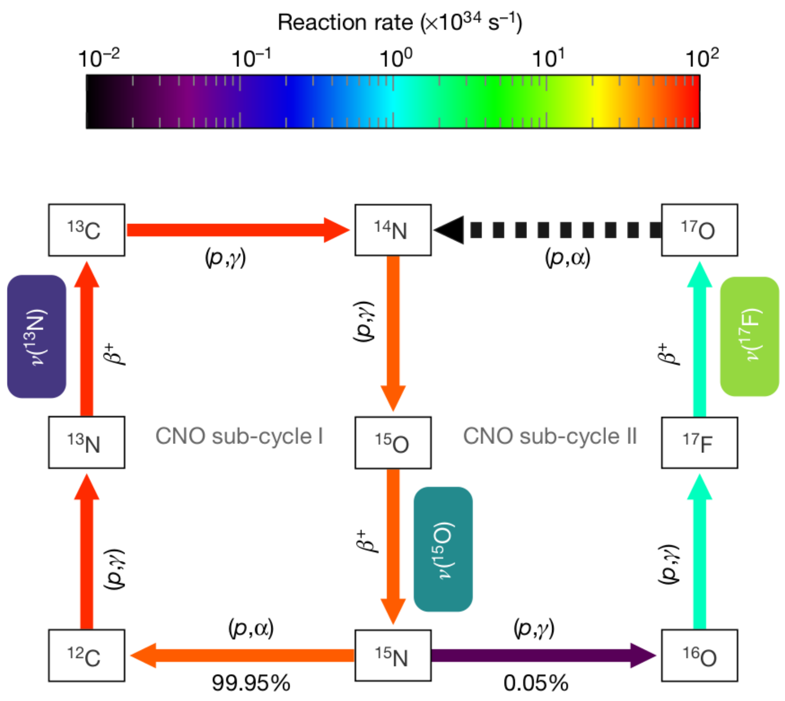
- Spectroscopy of the *pp* chain neutrinos
- Observation of the CNO cycle neutrinos

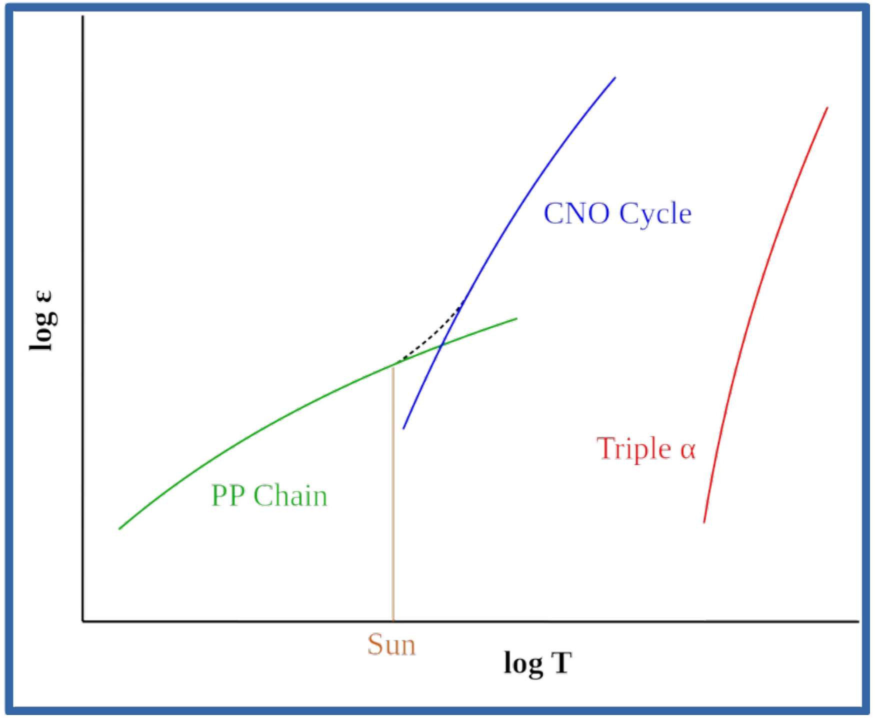
3. **SuperKamiokande**

- Oscillation physics with ^{8}B solar neutrinos

PP vs CNO Competition







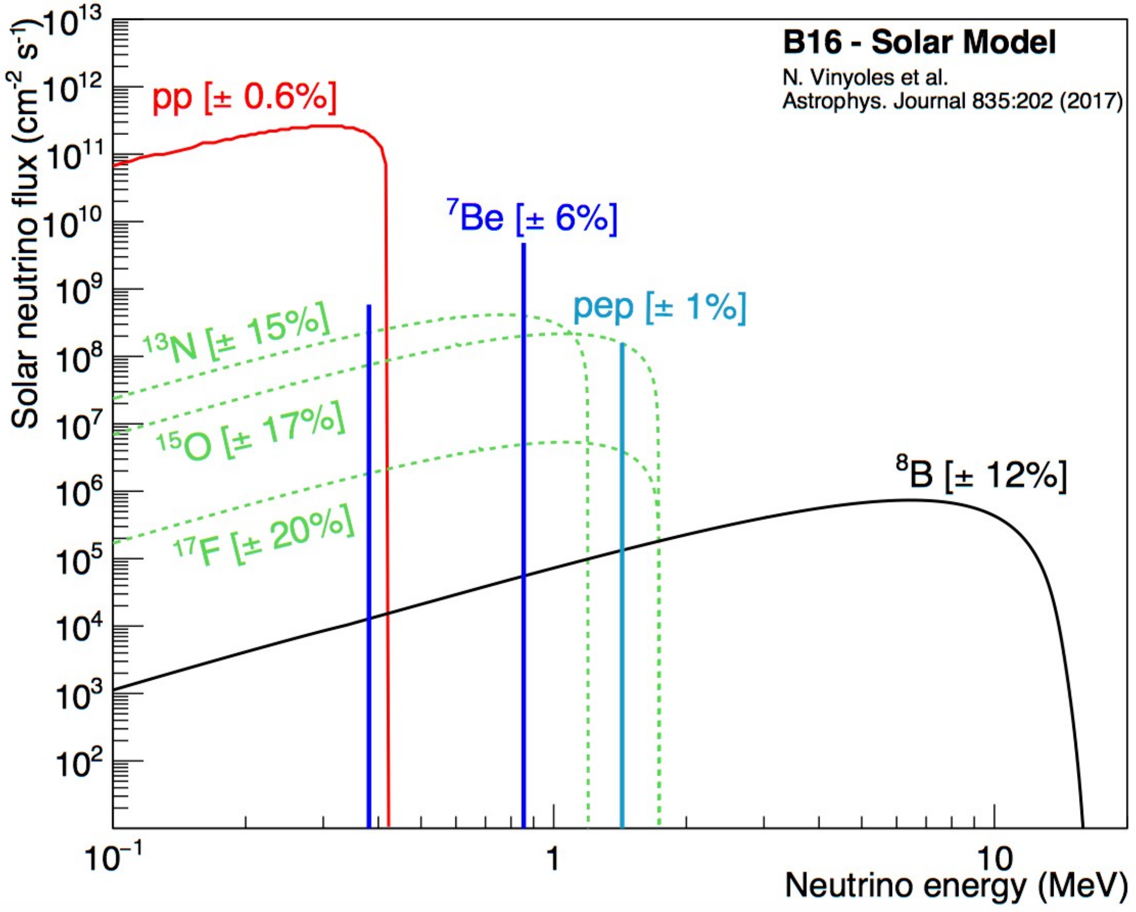
- Intense neutrinos from nuclear fusion in the Sun's core
- Majority (99%) from pp-chain with subdominant contribution from CNO cycle
- What's left in solar neutrinos?
    - Help understanding solar interior (metallicity problem)
- **Precision test of the MSW oscillation model**
    - Precise measurement of spectrum at the vacuum-to-matter transition region
    - Measurement of Day/Night asymmetry

The CNO cycle dominates in stars heavier than 1.3 M

Solar neutrinos as sensitive tool to test solar models: expected fluxes

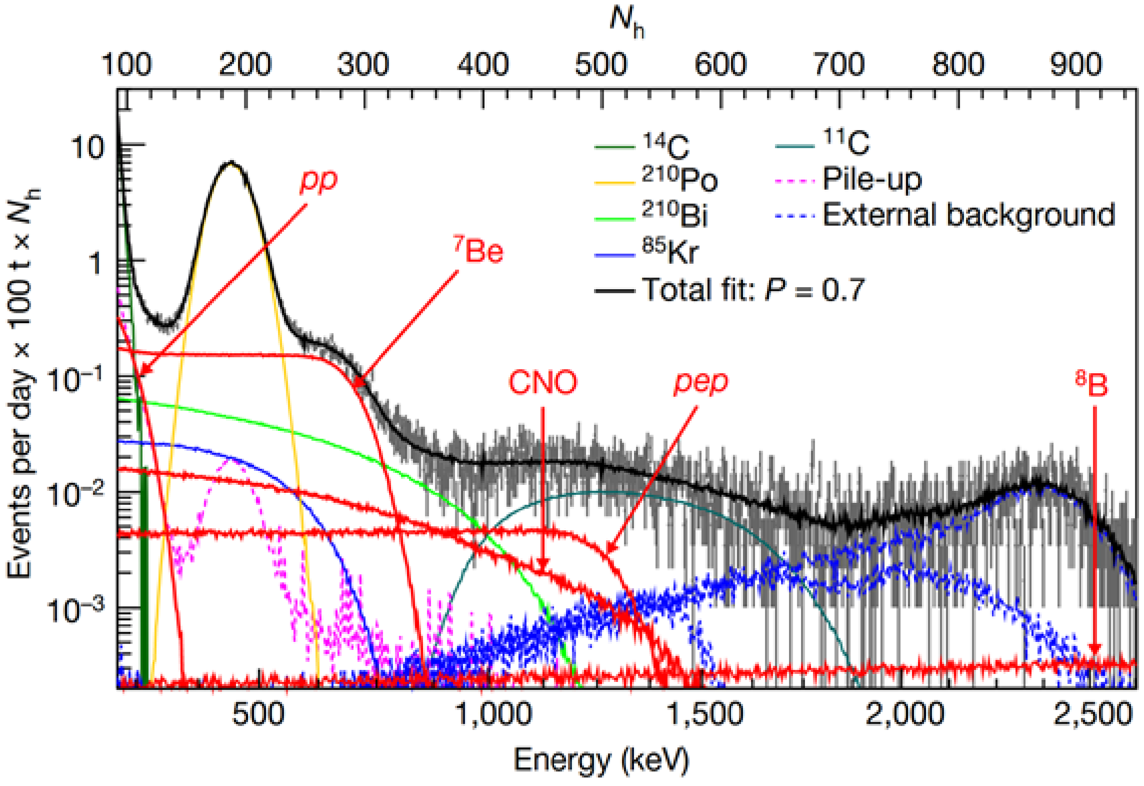
FLUX	B16-GS98
pp (10^{10} cm^{-2} s^{-1})	5.98(1±0.006)
pep (10^{8} cm^{-2} s^{-1})	1.44(1±0.01)
^{7}Be (10^{9} cm^{-2} s^{-1})	4.94(1±0.06)
^{8}B (10^{6} cm^{-2} s^{-1})	5.46(1±0.12)
^{13}N (10^{8} cm^{-2} s^{-1})	2.78(1±0.15)
^{15}O (10^{8} cm^{-2} s^{-1})	2.05(1±0.17)
^{17}F(10^{6} cm^{-2} s^{-1})	5.29(1±0.20)

N. Vinyoles et al.,
Astrophys. J. 835 (2017) 202



Original motivation of the first experiments on solar ν was to test the Standard Solar Model (SSM)

BOREXINO – real-time solar neutrino spectroscopy

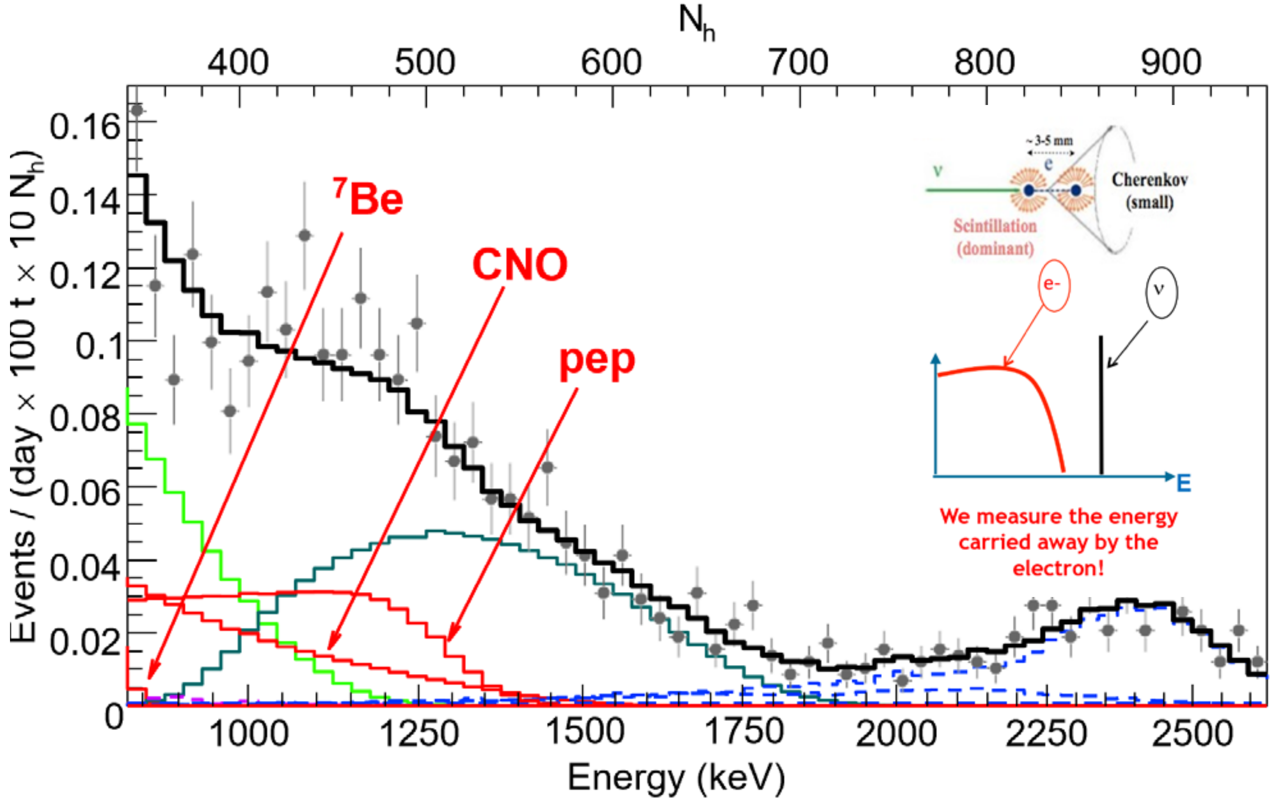


The Borexino Collaboration. *Comprehensive measurement of pp-chain solar neutrinos*. ***Nature*** **562,** 505–510 (2018)

From the measured interacton rates and assuming HZ-SSM fluxes we get electron neutrino survival probability from 60 keV to >10 MeV.

Selected the innermost β-like events

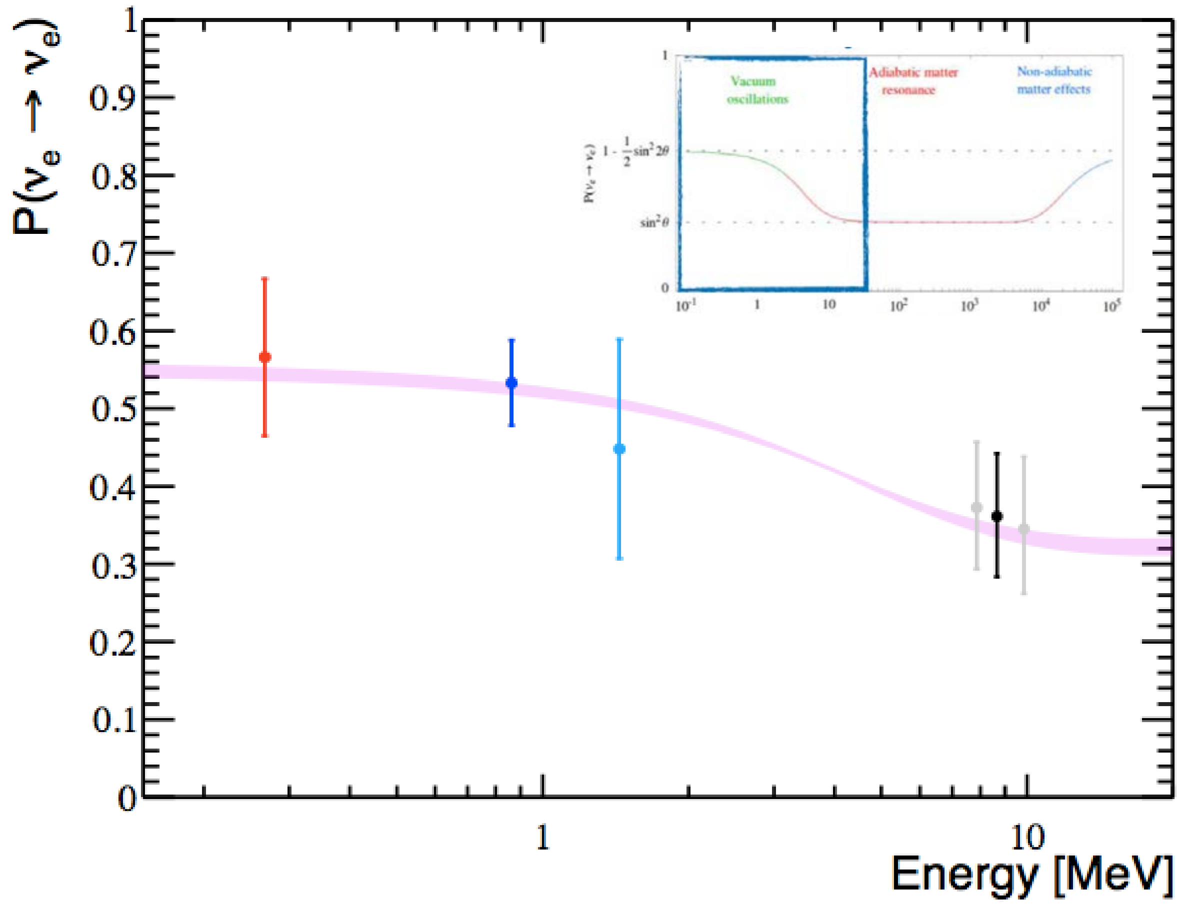
Radius <2.4 m Ps-LPR < 4.8



- $P_{ee}(pp)$= 0.57±0.10 $P_{ee}(^7Be, 862keV)$= 0.53±0.05
- $P_{ee}(pep)$= 0.43±0.11 $P_{ee}(^8B)$= 0.37$\underline{+}$0.08

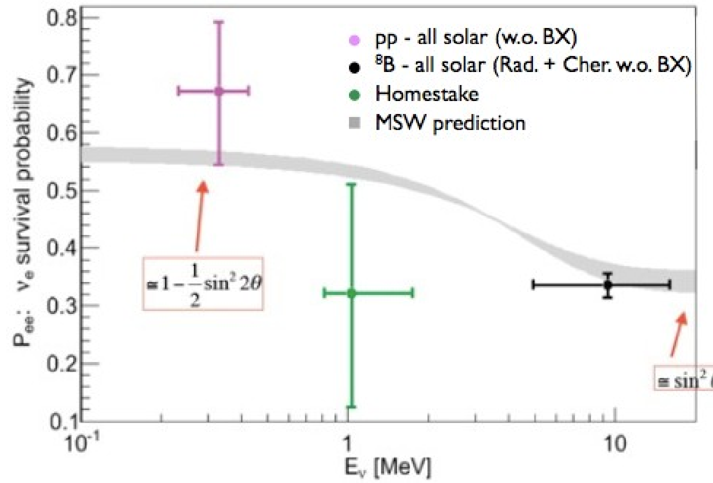
P_{ee}: Borexino impact

Borexino now



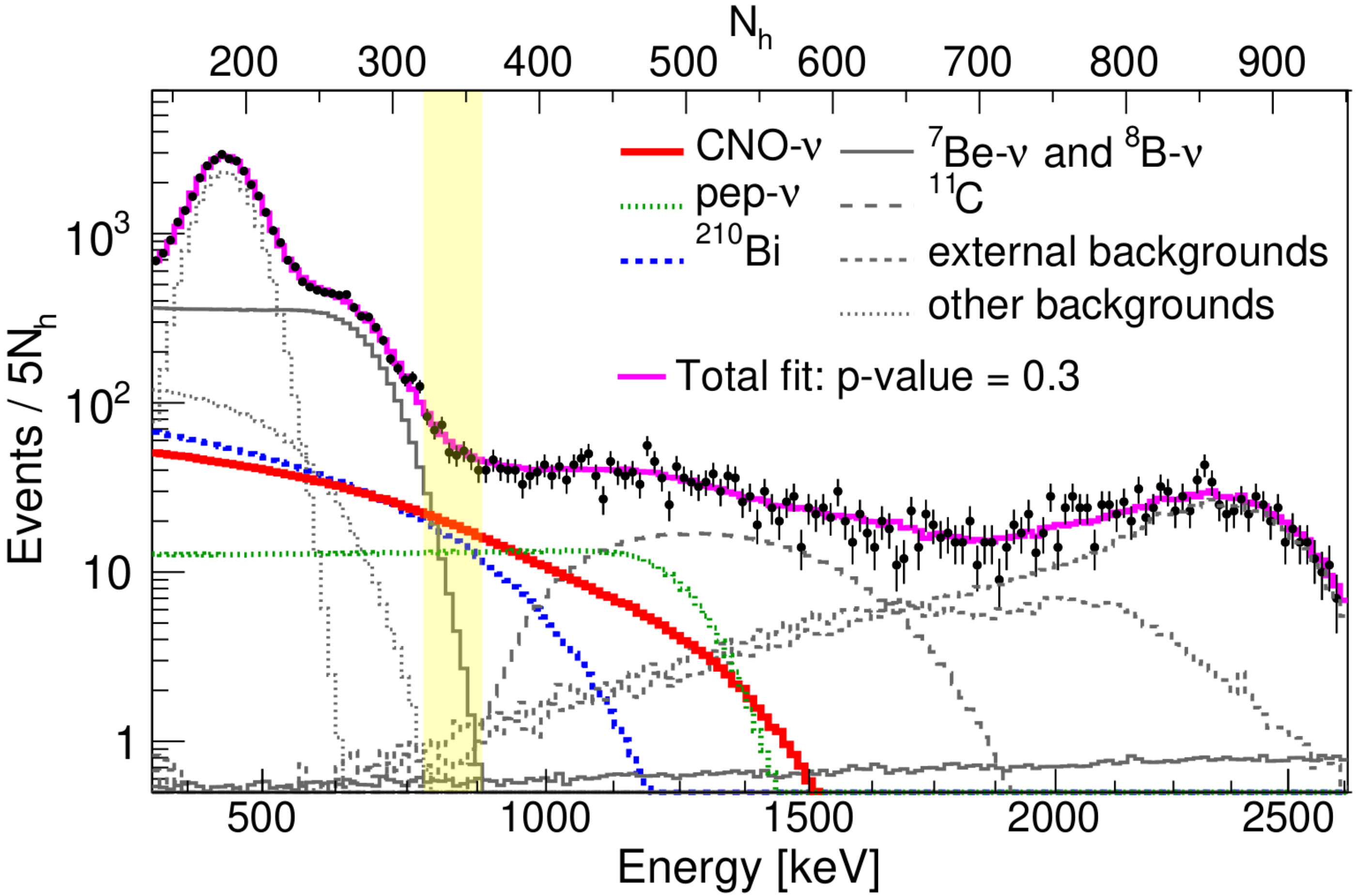
Pee - electron neutrino survival probability

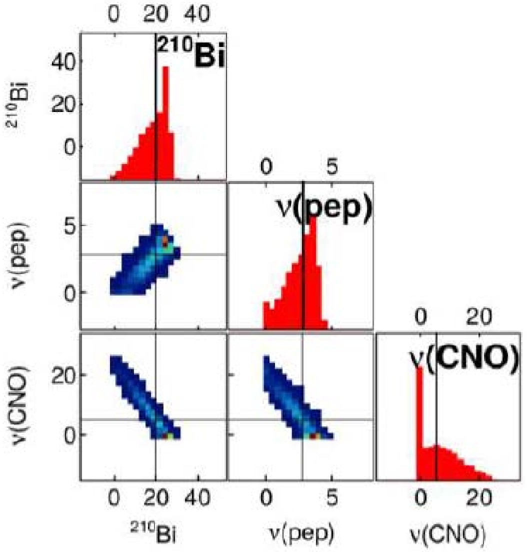
Before Borexino

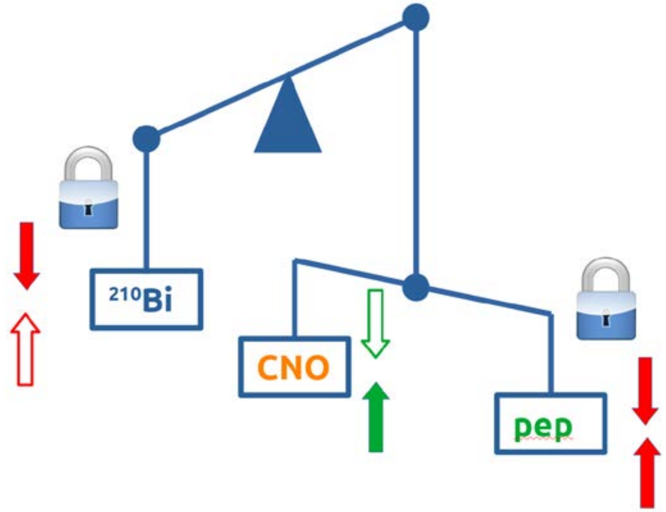


1. Borexino has measured the electron neutrino Pee in the *vacuum regime*, where, according to the MSW- LMA model, the vacuum dominates
2. The Borexino data allowed to probe the vacuum– matter transition from a single experiment.
3. Despite the uncertainty of the various points, that incorporate both the experimental errors and the SSM uncertainties, the experimental results seem in agreement with the predictions of the MSW-LMA model.

CNO - challenges



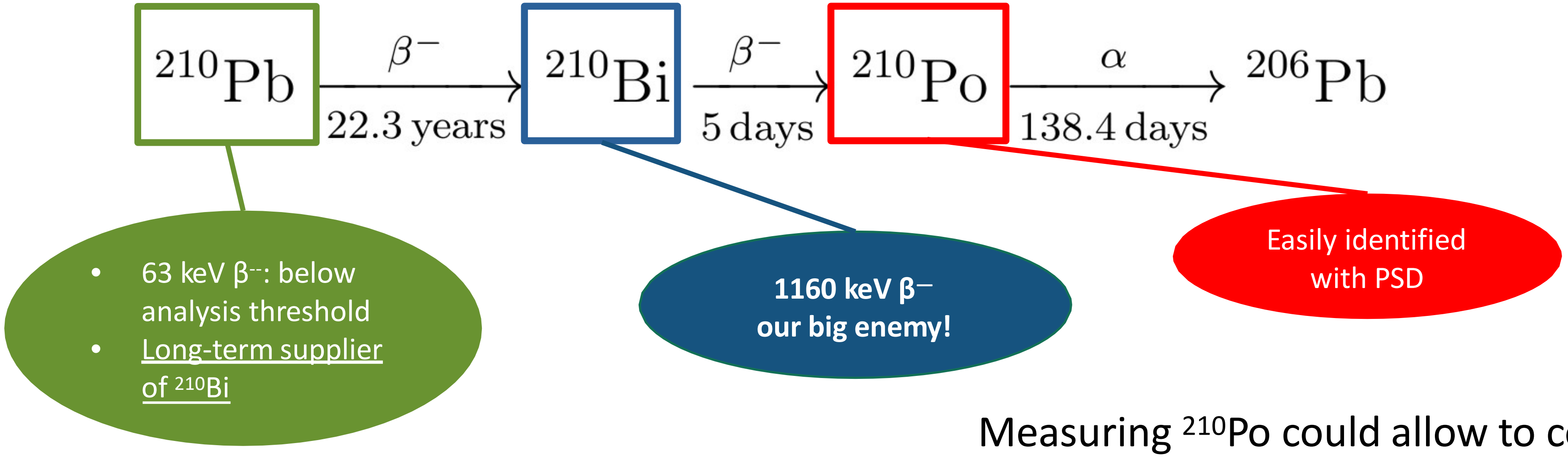




- *pep* rate: gaussian penalty at SSM prediction
- ^{210}Bi rate: semi-gaussian penalty at our upper limit

Strategy: independent constraint of pep and Bi-210

Strategy for ^{210}Bi constraint



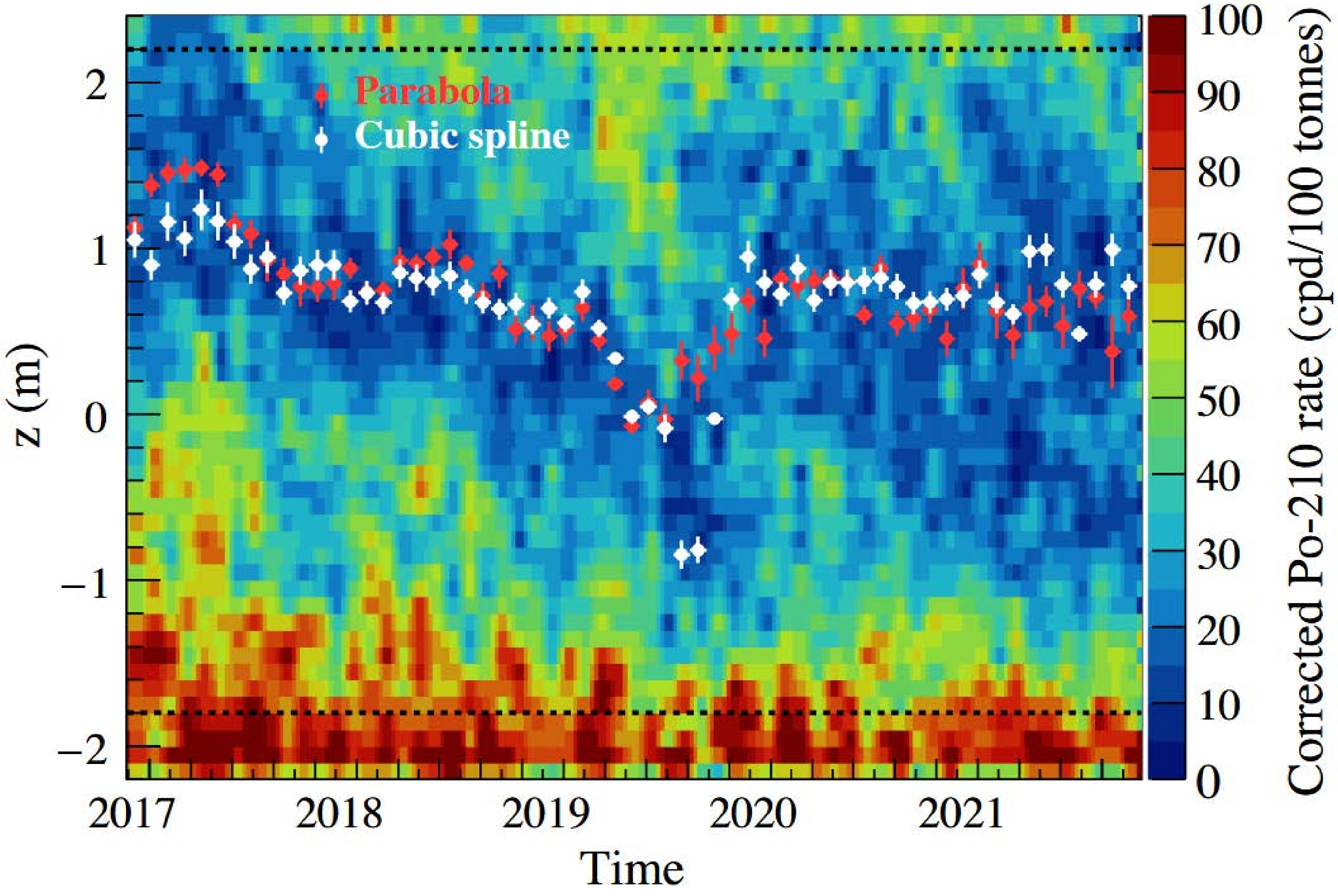
Measuring ^{210}Po could allow to constraint ^{210}Bi

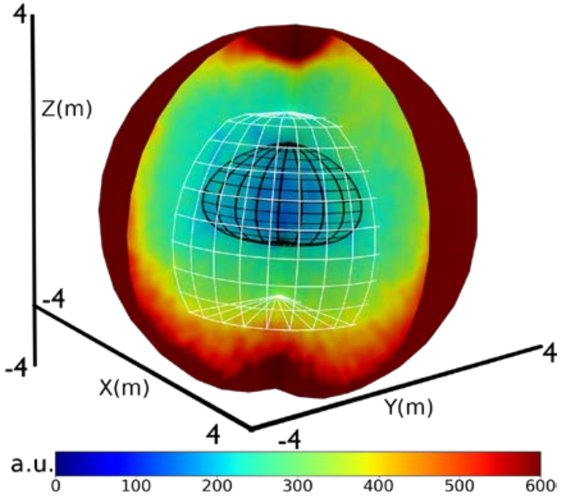
…

If only we had secular exquilibrium!

Constraining the decay rate of, a daughter of ^{210}Pb contaminating the scintillator, is a key requirement for the CNO analysis and is achieved by measuring the α decay rate of the ^{210}Bi daughter, ^{210}Po.

The Low Polonium Field





Clean region in the core of the detector is created: LPoF

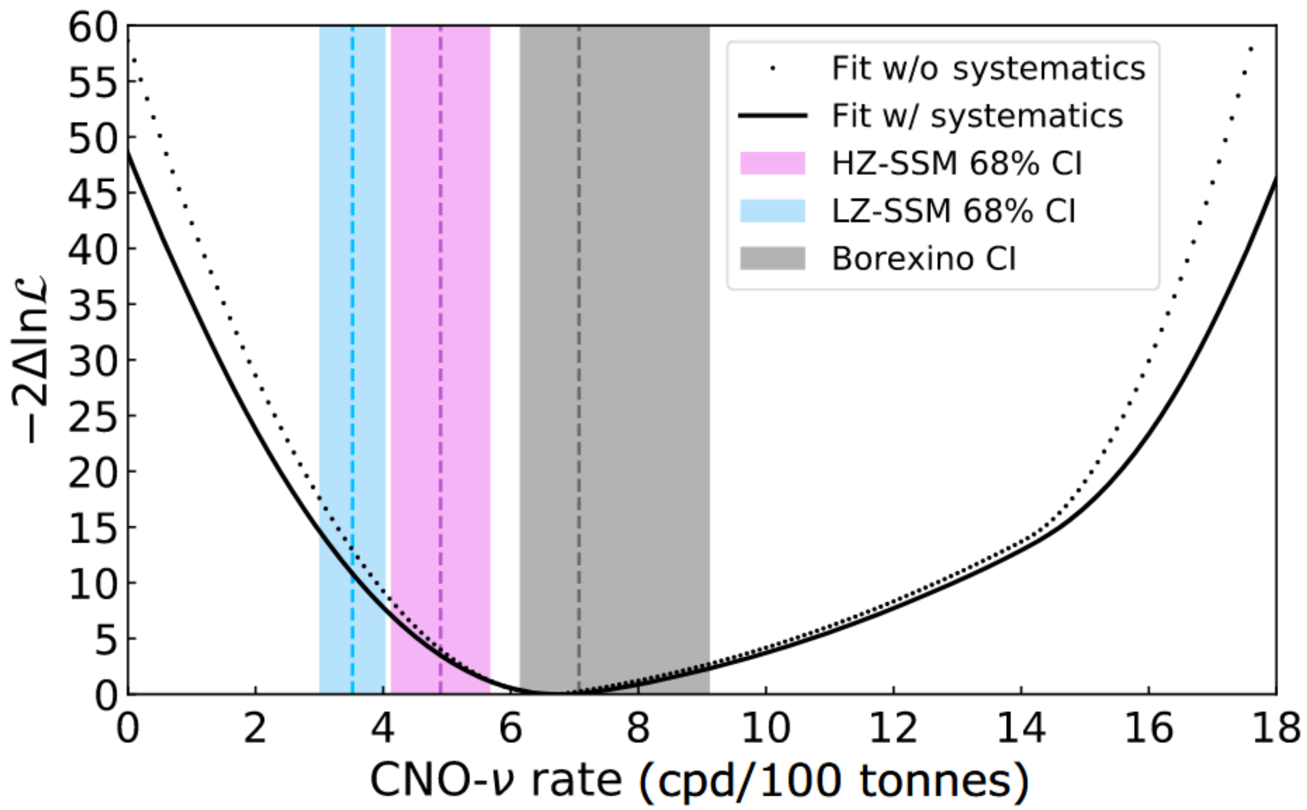
We extract the minimal ^{210}Po rate value, that is an upper limit on ^{210}Bi rate as a half-Gaussian constraint in the analysis

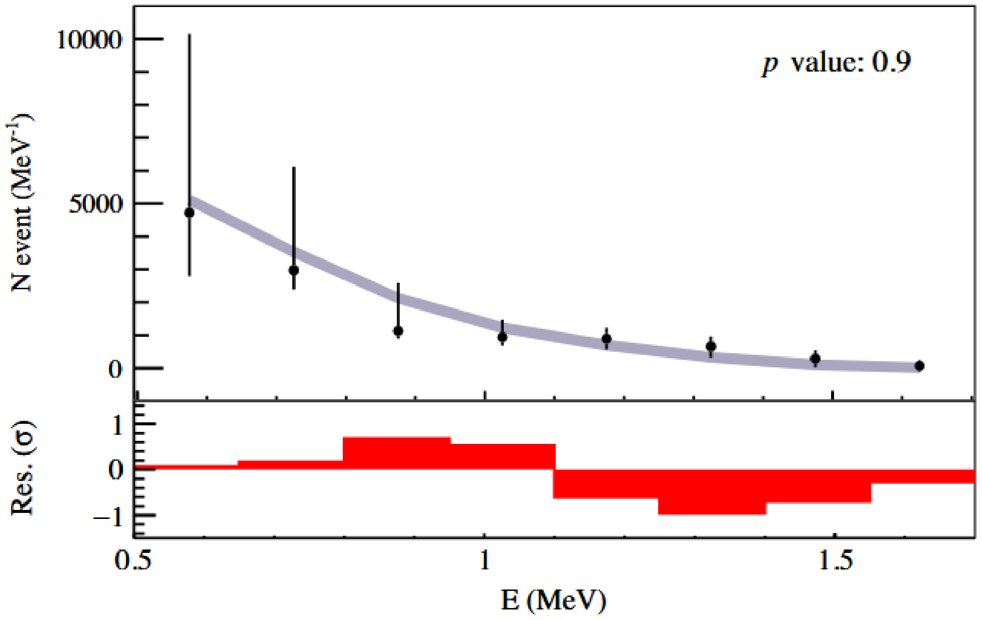
Bi < 10.8 +/- 1.0 cpd/100t (stat + sys)

$R(^{210}Po_{min}) = R(^{210}Bi) + R(^{210}Po_{vessel}) > R(^{210}Bi)$

CNO fit results

The Borexino Collaboration. *Improved Measurement of Solar Neutrinos from the Carbon-Nitrogen-Oxygen Cycle by Borexino and Its Implications for the Standard Solar Model.* arXiv:2205.15975





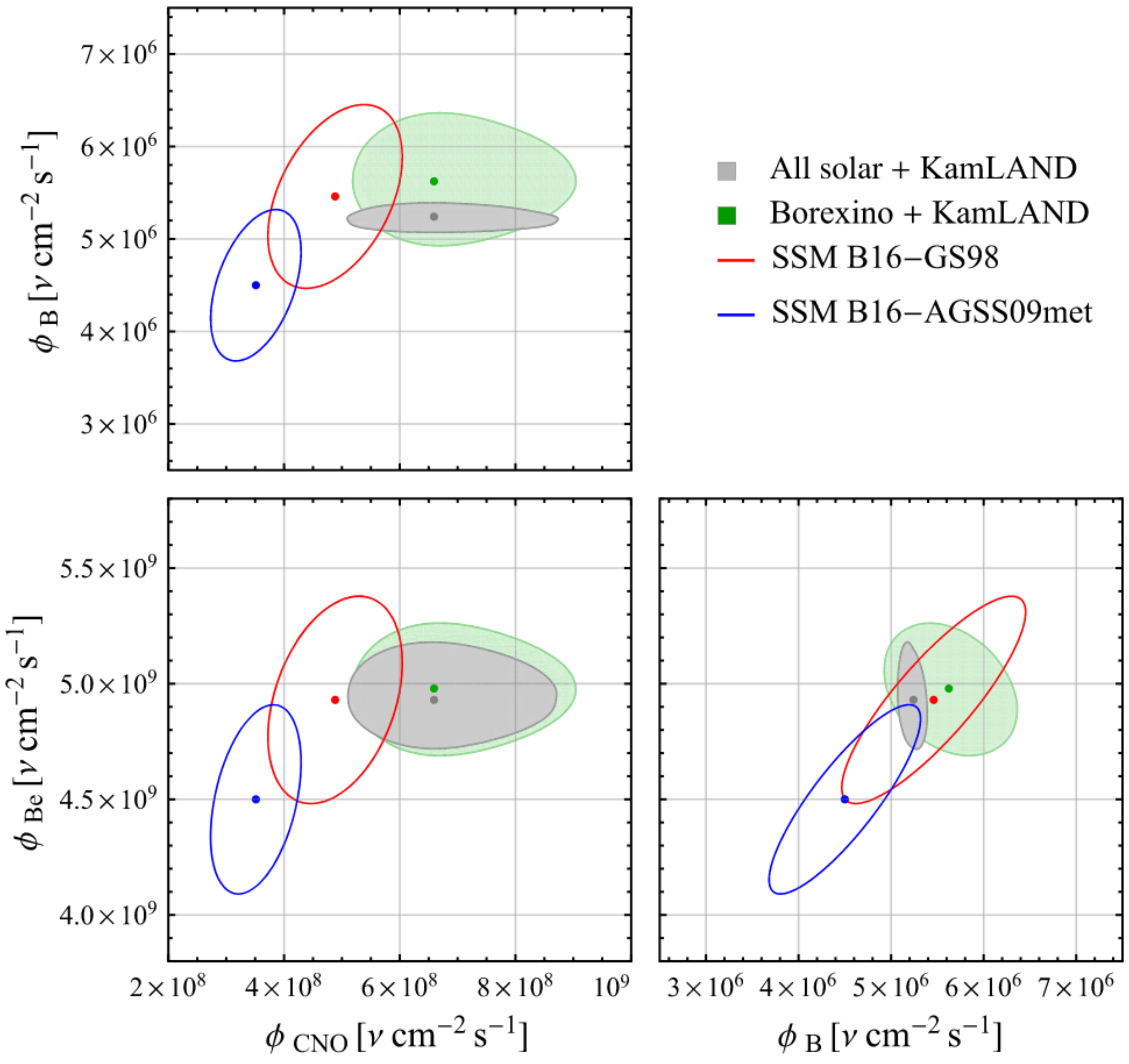
Spectral shape of the events after subtraction of all known backgrounds.

Result (68% CL stat + sys) = R_{CNO} = $6.7^{+2.0}_{-0.8}$ cpd/100 t

Φ(CNO with sys) = $6.6^{+2.0}_{-0.9} \times 10^8$ cm^{-2} s^{-1}

The result excludes the no-CNO-signal hypothesis at about 7σ C.L.

Solar implications: Global Analysis



2-dimensional planes C.I. for ν fluxes: ^{7}Be-ν, ^{8}B-ν,CNO

General agreement with SSM-HZ scenario

Binary hypothesis test: HZ vs LZ

→ Assuming SSM-HZ, Borexino results (7Be-ν + 8B-ν + CNO-ν), SSM-LZ is disfavored at ~3.1σ level

The Borexino Collaboration. *Improved Measurement of Solar Neutrinos from the Carbon-Nitrogen-Oxygen Cycle by Borexino and Its Implications for the Standard Solar Model.* arXiv:2205.15975

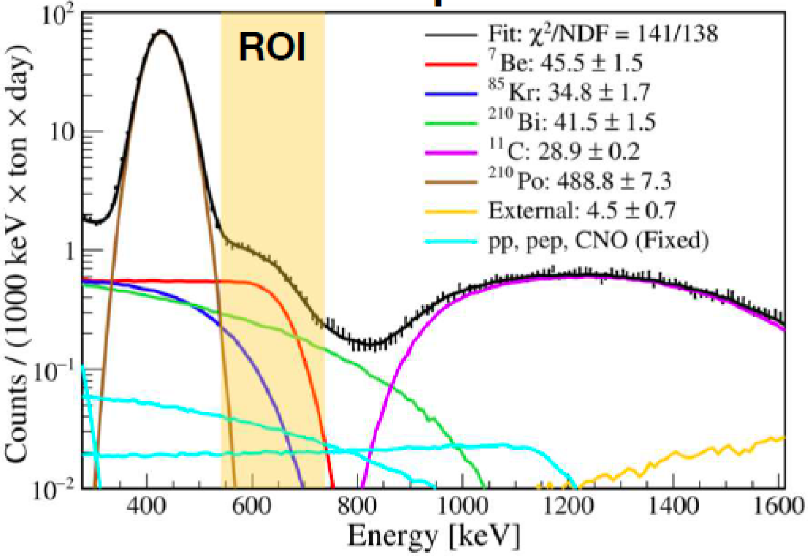
Future: First directional measurement of sub-MeV solar neutrinos

CID: Correlated and Integrated Directionality

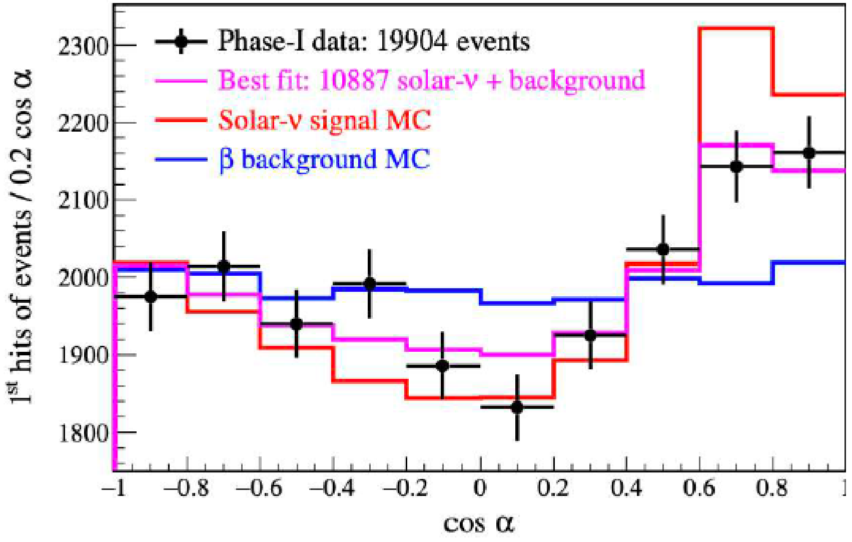
Exploit 1st and 2nd hit of each event (characterized by a high fraction of Cherenkov light)

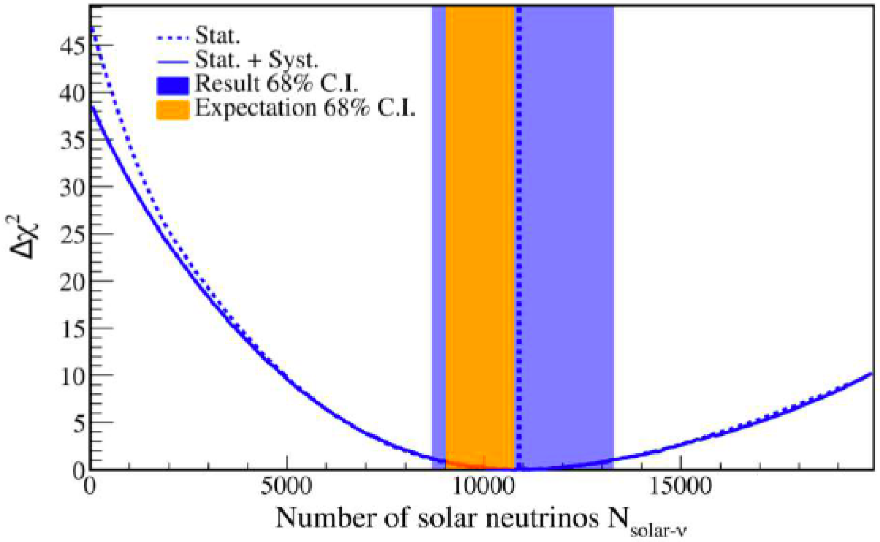
Results:

Phase-I spectral fit



cosα distribution





First directional measurement of sub-MeV solar neutrinos:

No-neutrino signal hypothesis ($N_{solar-\nu} = 0$) is **rejected with > 5σ**

$R(^7Be)_{CID} = 51.6^{+13.9}_{-12.5}$ cdp/100t (Compatible with SSM and spectral fit)

First Directional Measurement of sub-MeV Solar Neutrinos with Borexino, *Phys. Rev. Lett.* 128 (2022) 091803.

Correlated and Integrated Directionality for sub-MeV solar neutrinos in Borexino, *Phys. Rev. D* 105 (2022) 052002.

Borexino

1. Borexino has been the first experiment probing **sub-MeV neutrinos in real-time,** and is still now **the unique experiment** able to proceed with these studies.
2. Borexino has measured for **the first time all pp chain nuclear reactions producing neutrinos**, measuring, in particular, simultaneously the pp, ^{7}Be, and pep neutrino flux, ^{8}B neutrinos with a low threshold and probing hep neutrinos.
3. These results paved the way to actual breakthroughs not only on Solar physics, but also on neutrino physics. **The ν_e survival probability in the vacuum regime is measured** for the first time by Borexino and the **vacuum–matter transition** has been probed by a single experiment. In addition, a number of non-standard neutrino interactions has been studied by Borexino with world leading limits.

Borexino


European Physical Society
PRIZE

The 2021 Giuseppe and Vanna Cocconi Prize
for an outstanding contribution to Particle Astrophysics and Cosmology

is awarded to the

Borexino Collaboration

for their ground-breaking observation of solar neutrinos from the pp chain and CNO cycle that provided unique and comprehensive tests of the Sun as a nuclear fusion engine

Luc Bergé
President
European Physical Society

Thomas Gehrmann
Chair
EPS High Energy and Particle Physics Division

Mulhouse, France, 26 July 2021

4. The detection of the CNO cycle closes a long history, which began in the 30s of the last century, when Hans Bethe and Carl Friedrich von Weizsacker, independently, proposed that the fusion of hydrogen in stars could also be catalyzed by nuclei heavier than He. Then the theory of energy generation hypothesizes that the CNO would be the primary channel for hydrogen burning in stars more massive than the Sun , and it is in fact the primary channel for hydrogen burning in the Universe. This hypothesis never received an observational confirmation until now, when Borexino **has observed CNO neutrinos** proving also that its contribution in the Sun is of the order of 1%.
5. First directional measurement of sub-MeV solar neutrinos.
6. When all solar neutrino fluxes measured by Borexino, including CNO, are combined, the LZ hypothesis is **disfavored at a level of 3.1σ.**
7. Again, thanks to the low intrinsic background, Borexino has **observed geo-neutrinos** with 5σ statistical significance and studied them to obtain Earth geo-physical and geo-chemical information.

SuperK – ^{8}B Solar neutrino analysis improvements

- **Detector simulation improvements**
  - Improved PMT hit timing simulation
  - Improved modeling of water quality non-uniformity
- **Analysis improvements**
  - Correction for PMT gain drift
  - Improved correction for non-uniform energy response

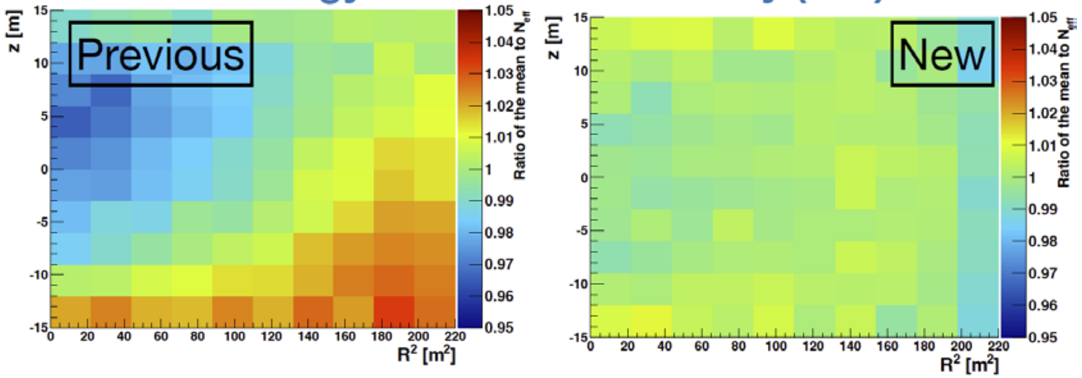
E-scale non-uniformity (MC) 1.7% → 0.5%

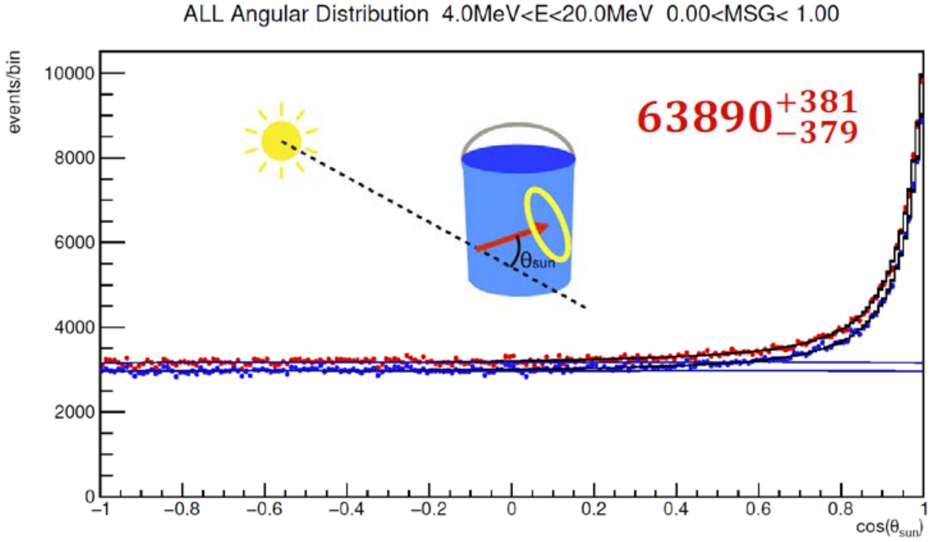
- **Improved spallation cut**
  - 12% more signal efficiency while keeping spallation rejection efficiency at a similar level (~90%)

Gained ~1 year worth statistics

Yasuhiro Nakajima *Recent results and future prospects from Super-Kamiokande* Neutrino 2020

Energy scale non-uniformity (MC)

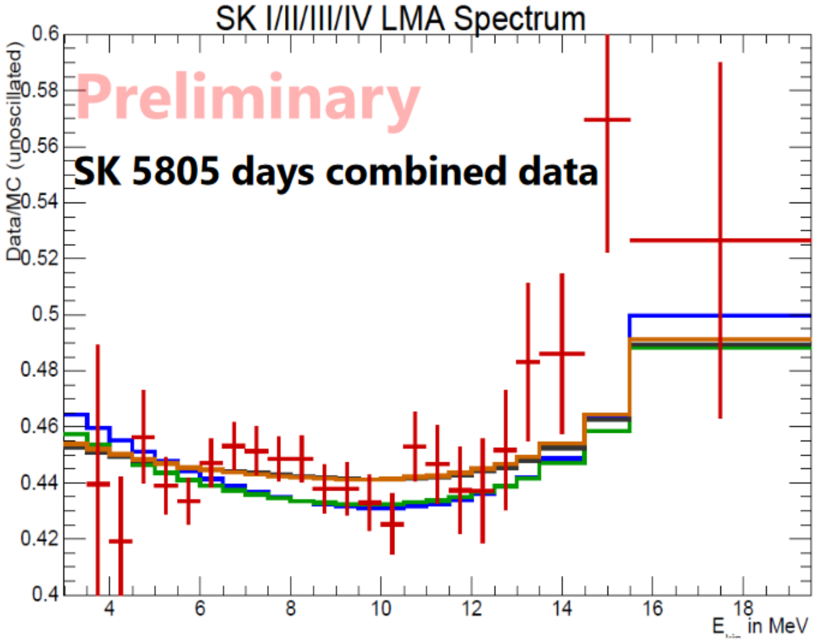




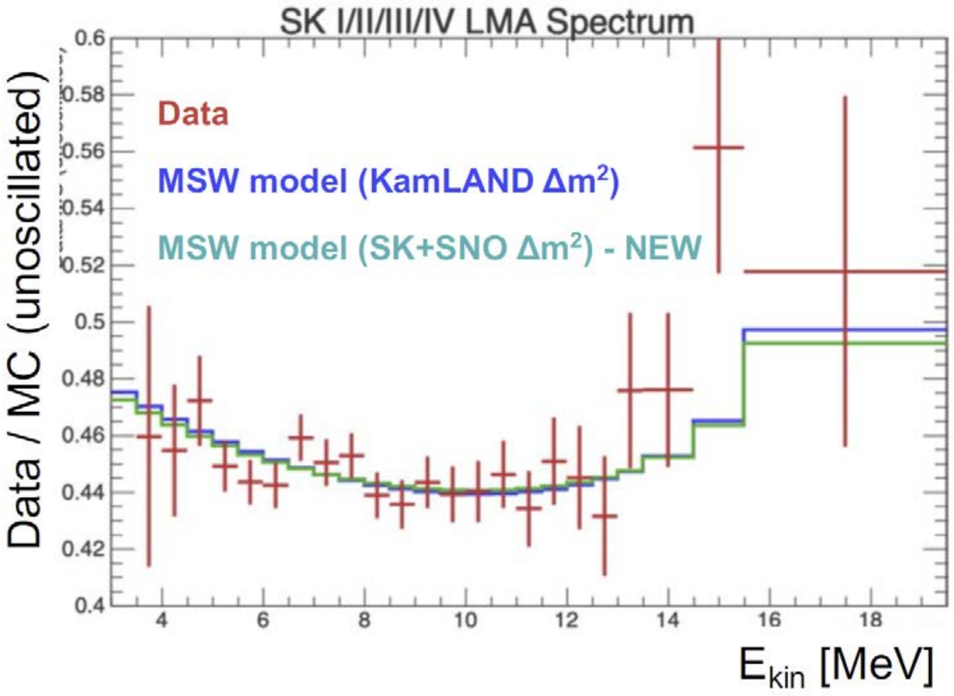
Locke, S. M. (2020). New Methods to Reduce Cosmogenic Backgrounds of Super-Kamiokande in the Solar Neutrino Energy Regime. *UC Irvine*.

New spectrum and Day/Nigth asymemetry measurements to test MSW

- Energy dependent survival probability Pee



TAUP 2019 - Yuuki Nakano and for the Super-Kamiokande collaboration 2020 J. Phys.: Conf. Ser. 1468 012189



TAUP 2021 – Livia Ludhova talk: Solar and Geoneutrinos

Data/MC ratio at E < 6 MeV slightly shifted upward

Shift of prediction due to improved detector simulation. Added statistics due to improved spallation cut.

Event migration due to new reconstruction

Day/Night asymmetry shift

Previous analysis used data up to Feb 2014 (SK-IV: 1664 days)

Added ~1200 days of data fluctuated towards smaller D/N asymmetry

Both impacted to the shift of best fit Δm^2_{21}

- Day/Night effect

$$A_{DN}^{Fit} = (-3.6 \pm 1.6(stat) \pm 0.6(syst))\,\% \rightarrow A_{DN}^{Fit} = (-2.1 \pm 1.1)\,\%$$

Neutrino 2020 Yasuhiro Nakajima *Recent results and future prospects from Super-Kamiokande*

SuperK - Oscillation Parameter Extraction

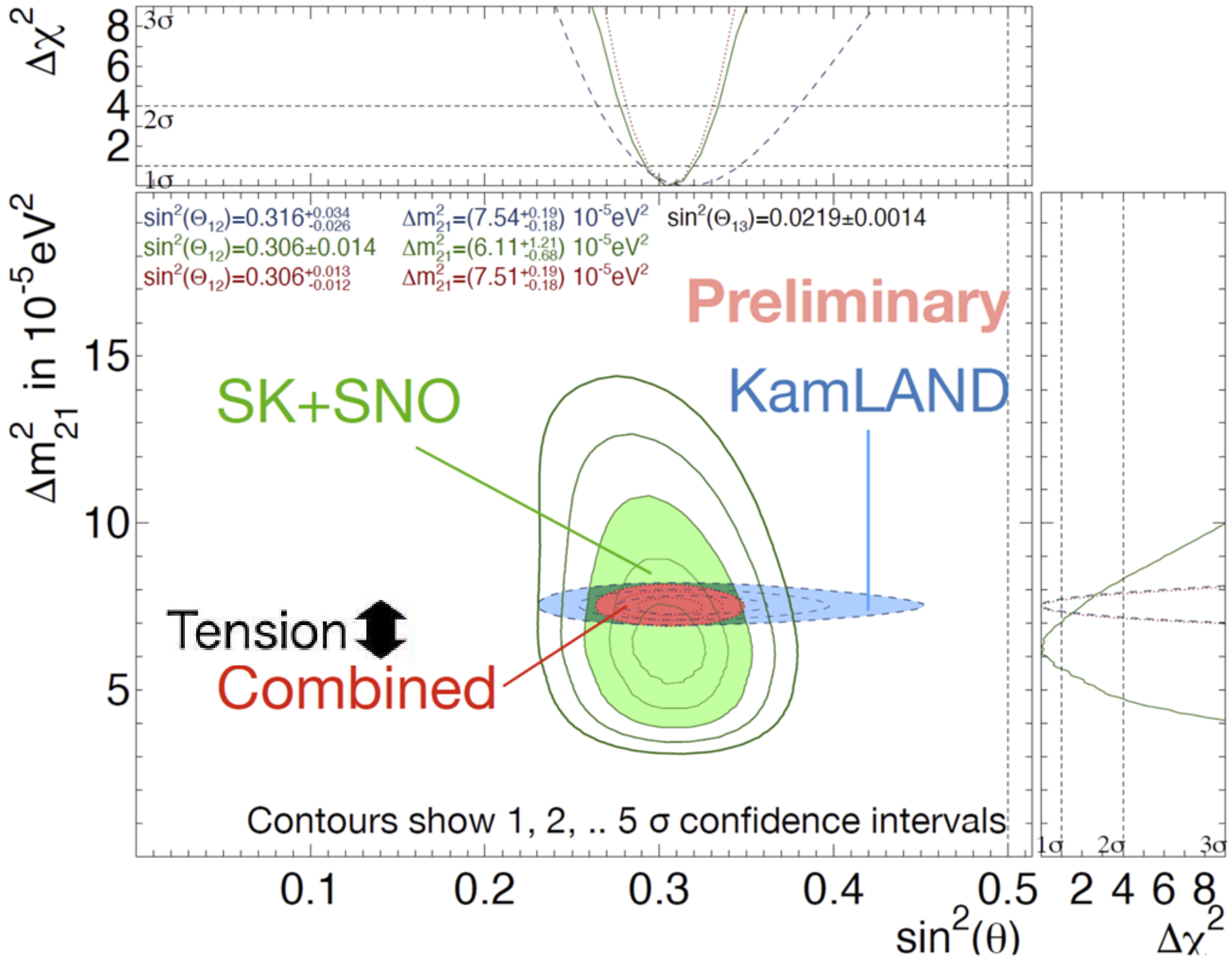
- use rate, spectral and day/night rate variation
- larger value of Δm^2 than before
- less tension (1.4 σ) with KamLAND (reactor antinu) -

- Oscillation parameters extracted by combining all SK data, as well as SNO and KamLAND data

	$\sin^2(\theta_{12})$	Δm^2_{21} [10^{-5} eV^2]
KamLAND	$0.316^{+0.034}_{-0.026}$	$7.54^{+0.19}_{-0.18}$
SK+SNO	0.306 ± 0.014	$6.11^{+1.21}_{-0.68}$
Combined	$0.306^{+0.013}_{-0.012}$	$7.51^{+0.19}_{-0.18}$

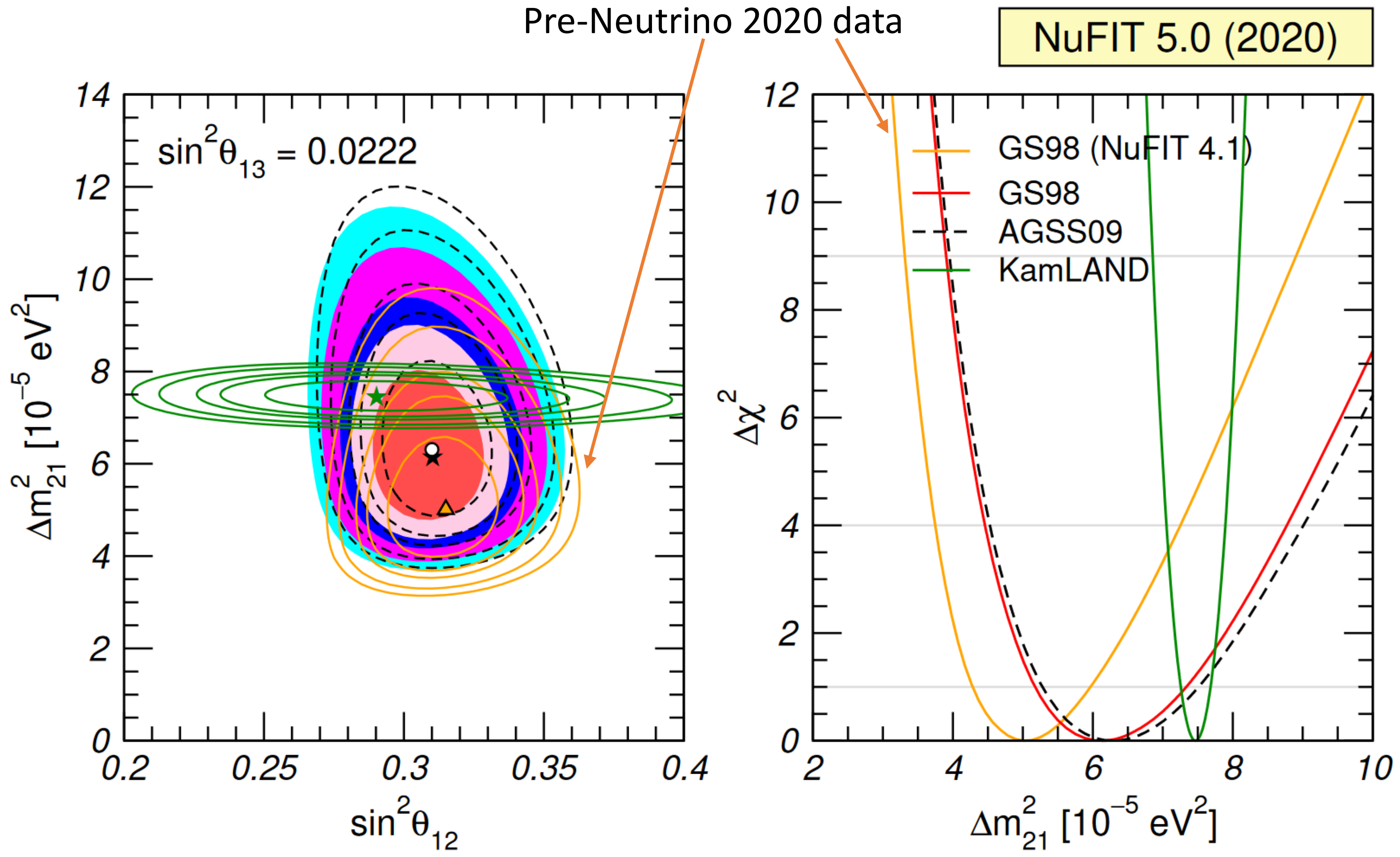
- Consistent θ_{12} values among experiments
- Solar best fit Δm^2_{21} lower than KamLAND, but difference is less than the previous analysis.

SK+SNO fit disfavors the KamLAND best fit value at ~1.4σ (was ~2σ)



Yasuhiro Nakajima *Recent results and future prospects from Super-Kamiokande* Neutrino 2020

Resolved tension in the solar sector



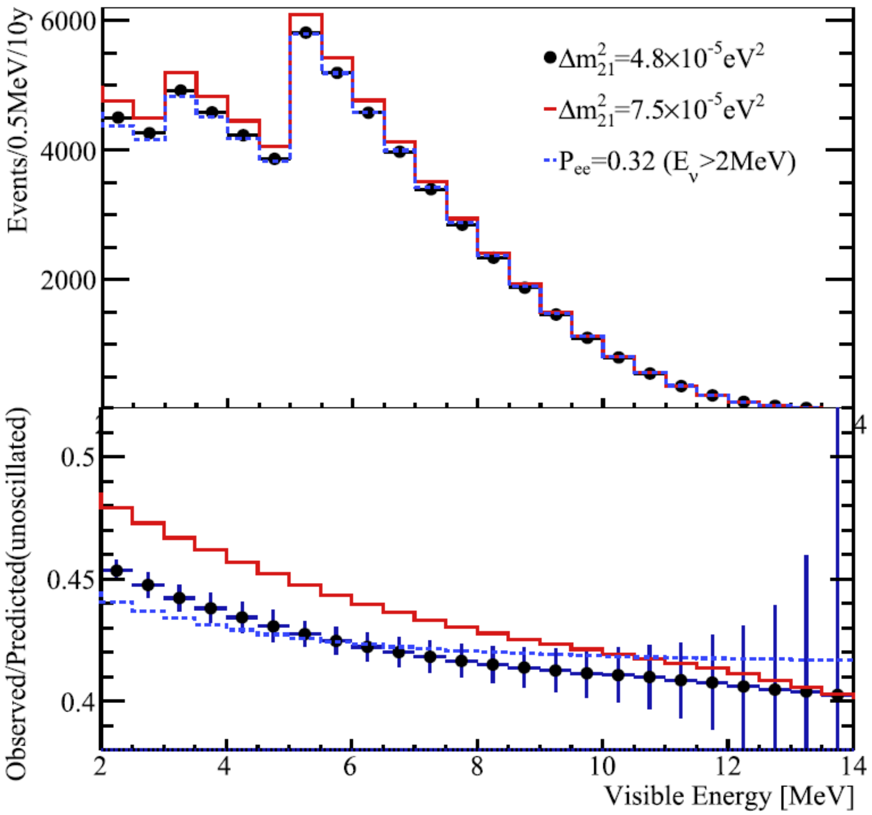
- With the new data the tension between the best fit Δm^2_{21} of KamLAND and that of the solar results has decreased.
- The best fit of KamLAND lies at 1.14σ in the analysis with the GS98 fluxes.
- This decrease in the tension is due to both, the smaller day-night asymmetry (and the slightly more pronounced turn-up in the low energy part of the spectrum which lowers it one extra unit.

Esteban, I., Gonzalez-Garcia, M., Maltoni, M. et al. *The fate of hints: updated global analysis of three-flavor neutrino oscillations.* J. High Energ. Phys. 2020, 178 (2020).

Future detectors

JUNO

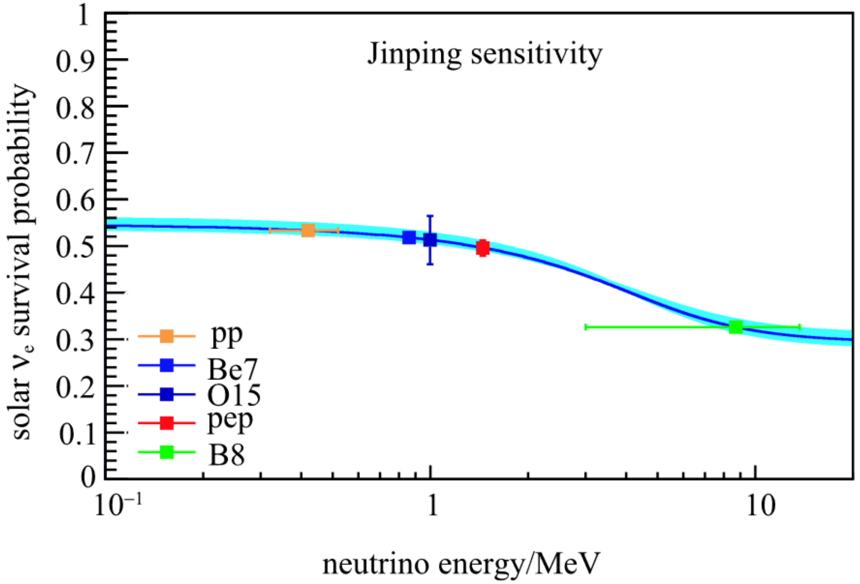
- 20 kt liquid scintillator
- excellent for B solar neutrino measurements,
- low-energy threshold,
- high energy resolution compared with water Cherenkov



Chinese Physics C 2021, Vol. 45 Issue(2) : 023004
DOI: 10.1088/1674-1137/abd92a

Jinping

- slow liquid scintillator
- total fiducial target mass of 2000 tons for solar neutrino



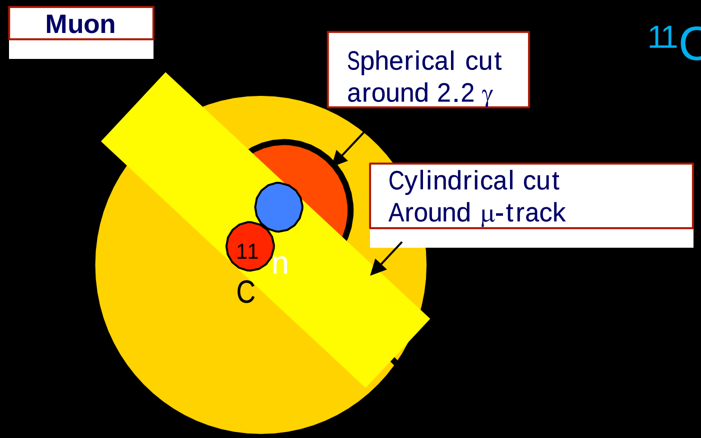
John F. Beacom *et al* 2017 *Chinese Phys. C* 41 023002

HyperKamiokande

- next generation large water Cherenkov detector
- water tanks provide the fiducial (total) volume of 0.19 (0.26) million metric tons

Backup slides

^{11}C cut: 1) the three fold coincidence



^{11}C is produced by muons together with neutron(s);

$\mu + {}^{12}C \;\; \text{->} \;\; {}^{11}C + n$

The likelihood that a certain event is ^{11}C is obtained using:

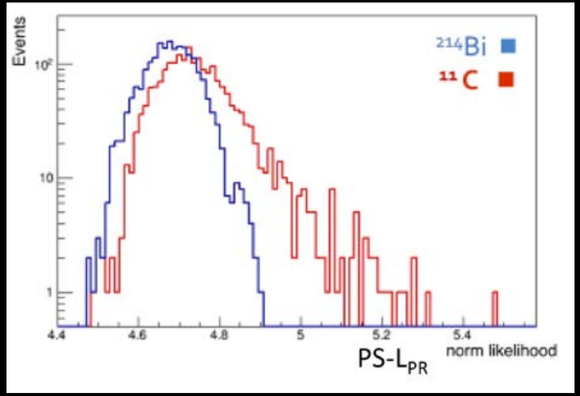
- Distance in space and time from the μ-track;
- Distance from the neutron;
- Neutron multiplicity;
- Muon dE/dx and number of muon clusters in an event;

The data-set is divided in two samples: one depleted in ^{11}C (TFC-subtracted) and one enriched in ^{11}C (TFC-tagged) which are simultaneously fit;

2) The β^+/β^- pulse-shape variable PS-L_{PR}:

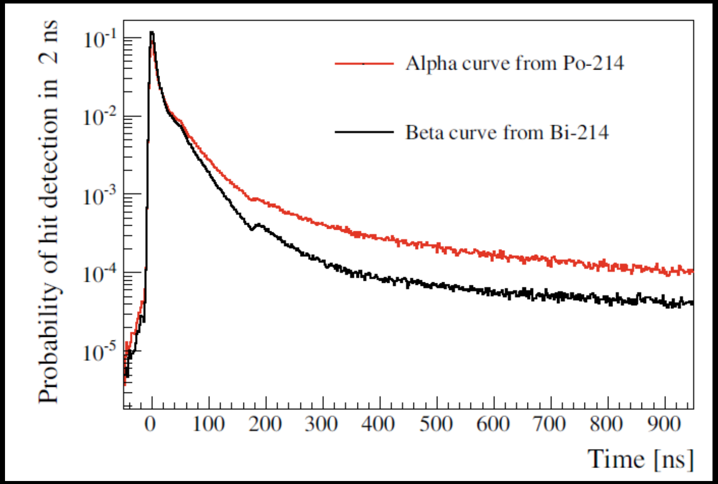
^{11}C decays β^+ : the probability density function (PDF) of the scintillation time profile is different for e^- and e^+ for two reasons:

- in 50% of the case e^+ annihilation is delayed by ortho-positronium formation (τ~3ns);
- e^+ energy deposit is not point-like because of the two annihilation gammas;



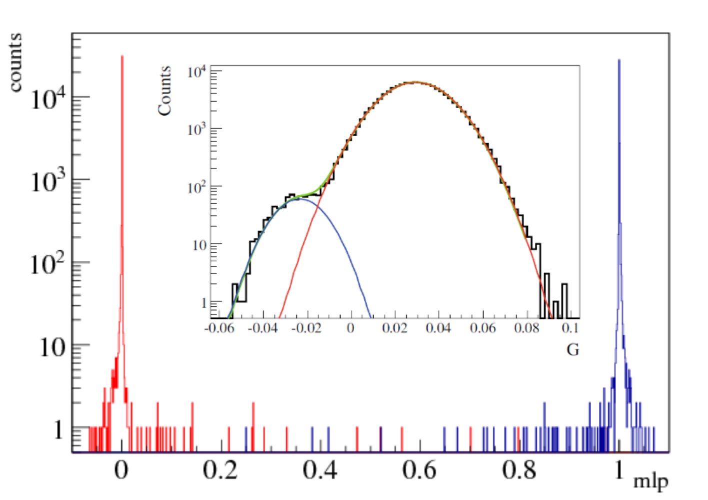
New discrimination parameter based on the output likelihood of the pos-reco alghoritm

α / β Discrimination with ANN



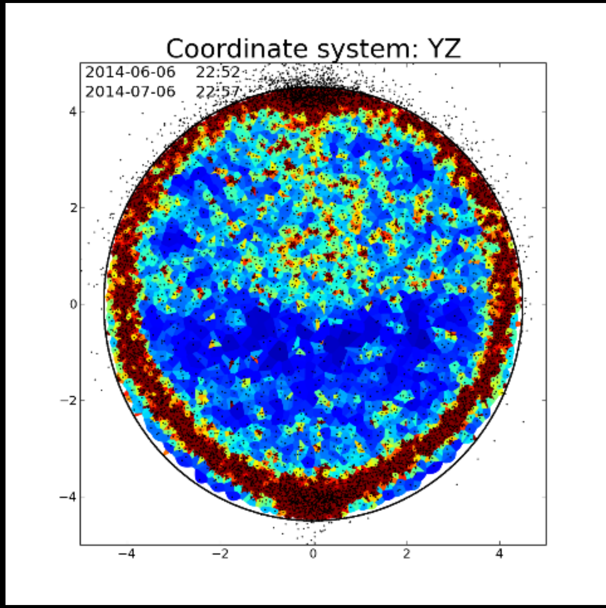
Solved problem of α discrimination from ^{210}Po (α)

Neural networks method



Galbiati, C., Misiaszek, M. & Rossi, N. Eur. Phys. J. A (2016) 52: 86.

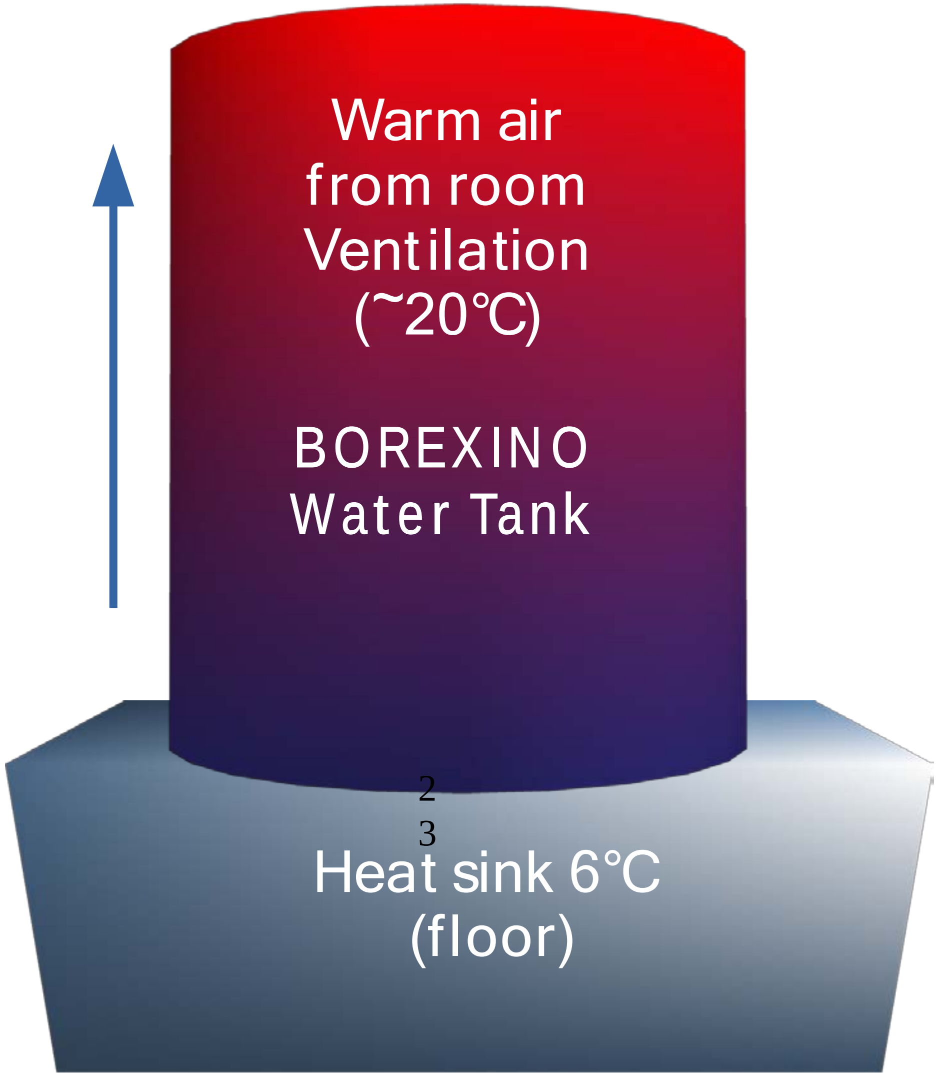
Example of α/β for ^{214}Bi-^{214}Po



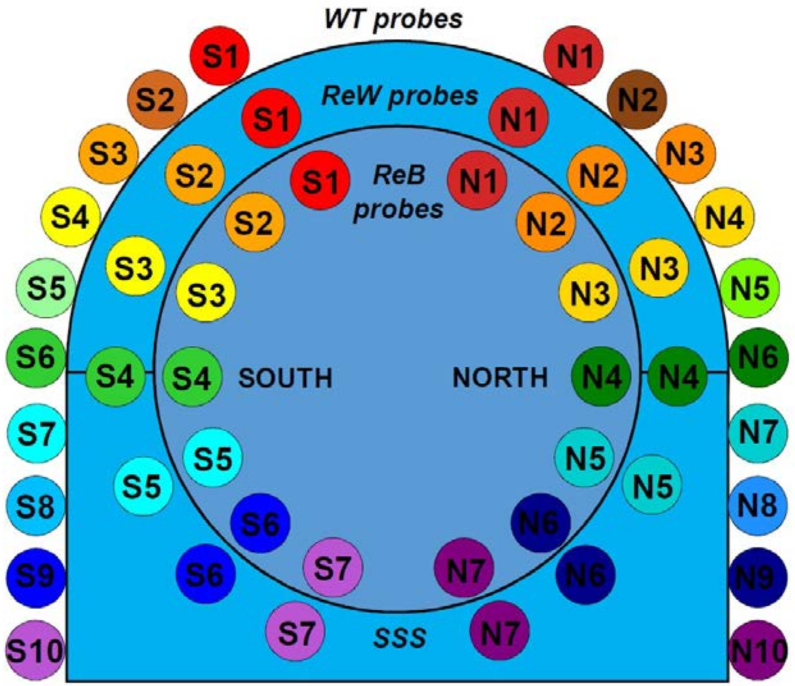
Thermal Insulation Program

Idea:
Strong and stable vertical gradient prevents convective motions

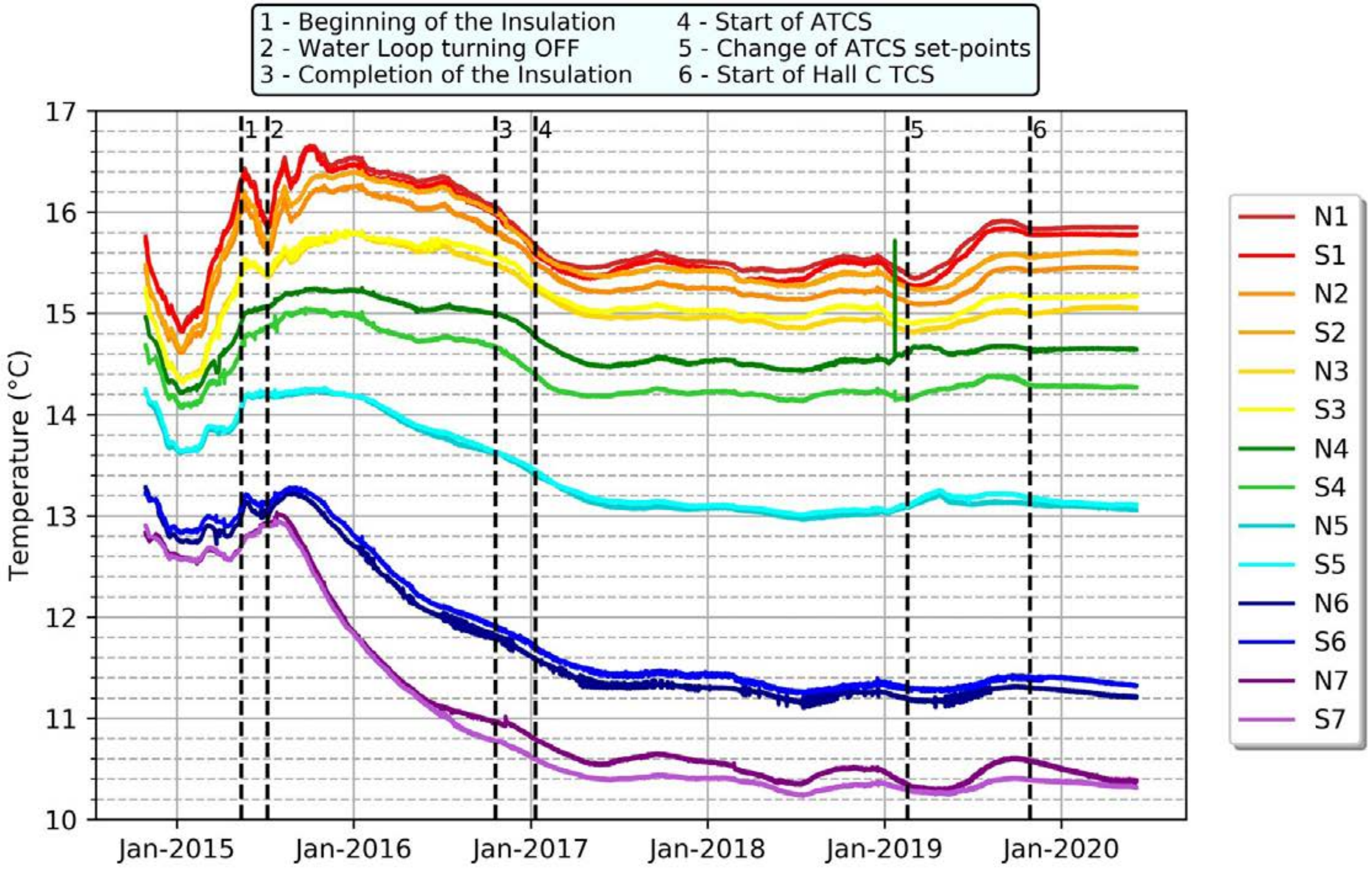
Milestones
2014: installation of temperature probes
Mid-2015: *beginning* of the insulation program
Late 2015: turning off of the *water recirculation system* in the water tank;
2016: first operation of the *active temperature control system* (ATCS)
Early 2019: change of the active control *set point*
Late 2019: installation and commissioning of the *hall C temperature control system.*



Effects on the temperatures



Temperature probes



Probes closer to the inner detectors with thermal program milestones